

The Murchison Widefield Array

Daniel A. Mitchell^{*}, Lincoln J. Greenhill, Stephen M. Ord and Gianni Bernardi

Harvard-Smithsonian Center for Astrophysics, e-mail: dmitchell@cfa.harvard.edu

Randall B. Wayth

Curtin University of Technology

Richard G. Edgar, Michael A. Clark, Kevin Dale and Hanspeter Pfister

Harvard University

Stewart J. Gleadow

University of Melbourne

W. Arcus, F.H. Briggs, L. Benkevitch, J.D. Bowman, J.D. Bunton, S. Burns, R.J. Cappallo, B.E. Corey, A. de Oliveira-Costa, L. Desouza, S.S. Doeleman, M.F. Derome, D. Emrich, M. Glossop, R. Goeke, M.R. Gopala Krishna, B. Hazelton, D.E. Herne, J.N. Hewitt, P.A. Kamini, D.L. Kaplan, J.C. Kasper, B.B. Kincaid, J. Kocz, E. Kowald, E. Kratzenberg, D. Kumar, C.J. Lonsdale, M.J. Lynch, S. Madhavi, M. Matejek, S.R. McWhirter, M.F. Morales, E. Morgan, D. Oberoi, J. Pathikulangara, T. Prabu, A. Rogers, J.E. Salah, R.J. Sault, N. Udaya Shankar, K.S. Srivani, J. Stevens, S.J. Tingay, A. Vaccarella, M. Waterson, R.L. Webster, A.R. Whitney, A. Williams and C. Williams

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It is shown that the excellent Murchison Radio-astronomy Observatory site allows the Murchison Widefield Array to employ a simple RFI blanking scheme and still calibrate visibilities and form images in the FM radio band. The techniques described are running autonomously in our calibration and imaging software, which is currently being used to process an FM-band survey of the entire southern sky.

RFI mitigation workshop

29-31 March 2010

Groningen, the Netherlands

^{*}Speaker.

[†]This work uses data obtained from the Murchison Radio-astronomy Observatory. We acknowledge the Wajarri Yamatji people as the traditional owners of the Observatory site. Support came from the U.S. National Science Foundation (grants AST-0457585 and PHY-0835713), the Australian Research Council (grants LE0775621 and LE0882938), the U.S. Air Force Office of Scientific Research (grant FA9550-0510247), the Smithsonian Astrophysical Observatory, the MIT School of Science, the Raman Research Institute, the Australian National University, the iVEC Petabyte Data Store, the Initiative in Innovative Computing and NVIDIA sponsored Center for Excellence at Harvard, and the International Centre for Radio Astronomy Research, a Joint Venture of Curtin University of Technology and The University of Western Australia, funded by the Western Australian State government. The speaker would also like to thank the National Astronomy and Ionosphere Center and Cornell University for travel support under grant AST-1013766.

1. Introduction

The Murchison Widefield Array (MWA) is both a pathfinder to the Square Kilometre Array and a stand-alone instrument, designed for cutting-edge science between 80 and 300 MHz via real-time imaging. This frequency band is rich with scientific potential, key science drivers being detection of redshifted 21 cm emission from neutral Hydrogen during the Epoch of Reionization, detection of radio transients, and space weather monitoring. However, the band can be difficult to observe in due to high levels of RFI. To reduce the need for complicated interference mitigation strategies, a remote site has been chosen so that RFI-free parts of the spectrum are available across the entire frequency range. While the 1024-input correlator and high-throughput software backend make the MWA a state-of-the-art instrument for advanced cancellation techniques, simple blanking methods have the advantage of simple statistics and are well suited to the low spectral and temporal occupancy at the site. Maintaining simple statistics is particularly important in the FM radio band, where it is anticipated that part of the weak neutral Hydrogen signal emitted during the Epoch of Reionization lies.

In addition to the excellent site, several design features help the MWA confront other challenges faced by modern and future radio facilities. For example, restricting the maximum baseline to a few kilometres reduces ionospheric distortion models to two dimensions, and real-time snapshot imaging allows the removal of these and various wide-field distortions in the image plane. Real-time processing also enables the large visibility data rate that comes with 512 antennas. A consequence of real-time processing and high data rates is that RFI algorithms need to be robust and efficient. A description of one such approach follows.

The Murchison Radio-astronomy Observatory

The Murchison Radio-astronomy Observatory (MRO) is a protected site in remote Western Australia that has been set aside as a park for a number of next-generation radio facilities. It consists of a zone, with a radius of 100 to 150 km, in which radio emission is restricted. While this level of protection is imperative for detection of extremely weak target signals, terrestrial radio waves can travel large distances via reflections off aircraft or meteor trails, or by occasional refractive ducting, while signals from satellites and aircraft can be in the direct line of sight.

FM radio signals from communities outside the restricted zone are an example of signals that will at times be reflected into the MRO. However, removing a modest amount of contaminated data will preserve this important band. A neighbouring instrument, the Experiment to Detect the Global Epoch of Reionization Signature (EDGES), removed all detected FM interference by filtering 3 percent of their 130 MHz band for up to 15 percent of the time during a typical day, with statistics collected over approximately 80 days [1].

2. The Real-Time Calibration and Imaging System

The MWA will consist of 512 antenna tiles, each a 4x4 array of dual-polarised, vertical bowtie dipoles. It is currently in a final design phase, with the first 32 antennas – along with prototype beamformers, receivers, correlator boards and computing hardware – being regularly used in an

end-to-end pipeline. The primary aim of the 32-tile system is to test hardware and firmware before the build out of the full array, and to aid the design of real-time software systems.

With the challenging dynamic range goal of a million to one in mind, a comprehensive, wide-field, real-time calibration and imaging system has been designed (see [2] and [3]). While the bulk of the real-time system is running routinely on simulated data using clusters with substantial graphics-card acceleration [4], for the 32-tile array the design has been significantly simplified. In keeping with the series of papers that accompany [5], and if the visibilities are dominated by a strong, unresolved, calibrator source, one can write the measured coherency matrix (containing the four polarisation terms of the visibility), after passing through antennas j and k , as

$$\mathbf{V}_{jk} = \mathbf{J}_j (\mathbf{S} \exp\{-i\phi_{jk}\}) \mathbf{J}_k^\dagger + \mathbf{N}_{jk}, \quad (2.1)$$

where a dagger superscript denotes a conjugate transpose, ϕ_{jk} is the phase shift due to differences in the path length from the source to the two antennas, and each bold symbol in (2.1) represents a 2×2 matrix; \mathbf{S} is the coherency matrix of the calibrator signal, the Jones matrices, \mathbf{J}_j and \mathbf{J}_k , describe the response of the antennas, and \mathbf{N}_{jk} contains the covariance of the polarised components of the noise. This noise comes mainly from the receivers and from other radio sources and, in the absence of RFI and significant secondary sources, we might expect it to have a Gaussian distribution. This is discussed further in section 3.

With a model for the source, $\mathbf{S} \exp\{-i\phi_{jk}\}$, and noise with an approximately Gaussian distribution, we can estimate the Jones matrix for each of the M antennas towards the calibrator using a least-squares approach (any amplitude or phase errors in the source model will be absorbed into these estimates). Following the matrix formalism suggested in [6] and described in [2], we search for the matrices $\hat{\mathbf{J}}_j$ that minimise

$$\chi^2 = \sum_{j=1}^M \sum_{k, k \neq j}^M \sigma_{jk}^{-2} \left\| \mathbf{V}_{jk} \exp\{+i\phi_{jk}\} - \hat{\mathbf{J}}_j \mathbf{S} \hat{\mathbf{J}}_k^\dagger \right\|_F^2, \quad (2.2)$$

where, for arbitrary matrix \mathbf{A} , the squared Frobenius norm $\|\mathbf{A}\|_F^2 = \text{Tr}(\mathbf{A}\mathbf{A}^\dagger)$, and $\sigma_{jk}^2 = \langle \|\mathbf{N}_{jk}\|_F^2 \rangle$. The solutions of (2.2),

$$\hat{\mathbf{J}}_j = \left(\sum_{k, k \neq j}^M \sigma_{jk}^{-2} (\mathbf{V}_{jk} \exp\{+i\phi_{jk}\}) \hat{\mathbf{J}}_k \mathbf{S}^\dagger \right) \left(\sum_{k, k \neq j}^M \sigma_{jk}^{-2} \mathbf{S} \hat{\mathbf{J}}_k^\dagger \hat{\mathbf{J}}_k \mathbf{S}^\dagger \right)^{-1}, \quad (2.3)$$

can be estimated from an initial set of Jones matrices and iterated until they converge to an optimal set, which can then be tracked.

The MWA is a snapshot instrument that uses (2.1) and (2.3) to produce calibrated images every eight seconds for each frequency channel and polarisation.¹ The images are regridded in real time to remove ionospheric and wide-field distortions, then integrated for several minutes. Eventually (2.3) will be updated for many calibrators, to build up direction-dependent gain models, but for the 32-tile array it is assumed that the shape of the primary beam response of each antenna is known, so that the models can be fixed by a single calibrator. It is also assumed that the primary beam shapes are the same for each antenna, so that the fixed beams are equal. These issues are described

¹ A separate algorithm is used to estimate deviations in ϕ_{jk} arising from the ionosphere.

in more detail in [7], where the real-time system was applied to the 32-tile array to track fields and integrate for many hours over much of the band from 90 to 210 MHz.

2.1 Flagging in the Real-Time System

With all of the dimensions that can be used to describe a typical observation – time, frequency, polarisation, antenna space, visibility space, image space, etc. – RFI mitigation will in general be most effective if one can find a projection in which the interference has little or no overlap with the signals of interest. For RFI that is contained within small segments of time and frequency, these two dimensions are obvious places to define the segmentation, and traditional techniques for blanking samples in time and frequency offer close to complete removal of the RFI at the expense of losing a small amount of the signal of interest.

The calibration and imaging system has a separate processing pipeline for each frequency channel, and this coupled with the real-time nature of the instrument makes it well suited to time and frequency blanking. While there are a number of places to set flags that can be used throughout the system to blank visibilities, it is only being done at the moment at the start of the software system. There are two detectors in use; one generates statistics for each frequency channel using neighbouring time samples, the other generates statistics for each time step using neighbouring frequency samples. Both detectors search for outliers using quartiles, where departures from the median value by more than a set multiple of the interquartile range are flagged for blanking. The interquartile range is a robust estimator of signal variance and is relatively cheap to compute (see, for example, [8]). Each visibility has an associated non-negative inverse-variance weight, and flags are set by making these negative. The calibration and imaging routines ignore samples that have a negative weight, and accumulated quantities that are missing samples due to flagging can simply be down weighted.

Development is underway to combine these detectors with goodness-of-fit estimates for the relations given at the start of this section. Also, since images are formed every eight seconds, various image processing tools can be run before integration, as a final check of data quality.

3. 32-Tile Results

The accuracy of calibration relations (2.1) and (2.3) will in general be diminished by any RFI in the coherency matrices. In this section some examples will be given showing the effect of both RFI and blanking. Data come from 21 of the 32 antennas and 6.2 MHz of bandwidth centred at 98 MHz. The antenna tiles have been steered to a field with a right ascension of 21.2 hours and a declination of -25.4 degrees, close to transit. The dominant object in the field is extragalactic radio source PKS J2107-2526, which is 61 Jy at 80 MHz [9]. This is one of the calibration pointings from a survey of the entire southern sky, which is currently under construction.

In the absence of RFI, the validity of the model given in (2.1) may be verified by inspecting noise matrices. The matrices that minimise (2.2) can be used to remove the calibrator power from the measured coherency matrix (known as peeling [10]), and if the assumptions hold then the remaining noise term should have a Gaussian distribution. The noise distribution after peeling is shown in figure 1. In this paper the fundamental accuracy of the calibration process is not being considered, and it is adequate to note that the noise distributions qualitatively match Gaussians.

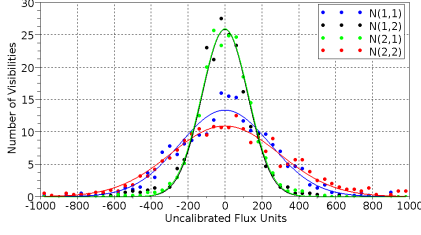


Figure 1: Histograms of the noise in each element of the coherency matrices – estimated by peeling the fitted source contribution – along with best-fit Gaussian curves. The peeling was done separately in each 40 kHz channel, and the statistics generated after combining the whole band (after dividing out a polynomial bandpass fit). Several histograms were averaged to reduce the scatter.

In figure 2 the amplitude of Jones matrix element $J_j(1,1)$ is shown as a function of time and frequency for one of the MWA antennas. To avoid known artifacts of the prototype receiving system, specific frequency channels were always blanked for these data, which can be seen as vertical bands of zero gain. About halfway through the scan RFI becomes evident and increases in power as the scan continues (the gain between FM radio frequencies is also affected by contamination of the bandpass fits).

It is clear that blanking has significantly reduced the RFI in these data. Jones matrix errors have been reduced by 30 dB, leaving gains that are stable in both frequency and time. During the worst of the RFI only about 12 percent, or ~ 0.76 MHz of bandwidth, was blanked. The quality of the unblanked data can be seen in figure 3, where images made with and without blanking are compared. The calibrator is completely buried in the presence of RFI, but can be imaged and peeled below the noise floor with the RFI flags.

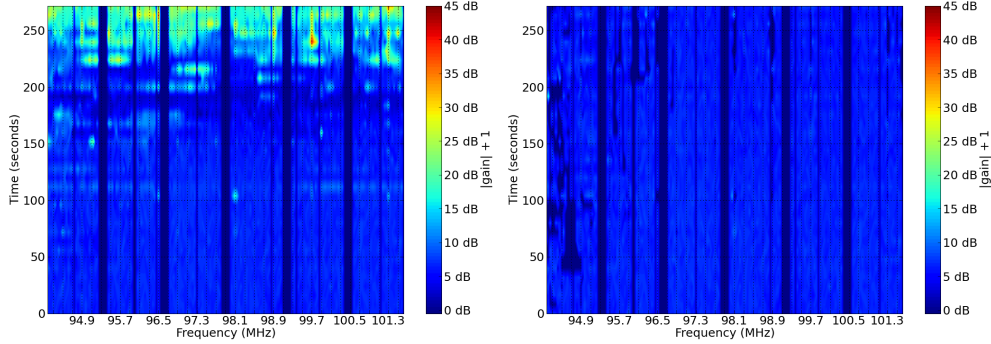


Figure 2: Jones matrix element $J_j(1,1)$ is shown as a function of time and frequency for MWA antenna tile 1, with dashed vertical lines indicating possible FM radio frequencies. The image on the left-hand side shows fitted gain amplitude without RFI blanking, while the image on the right-hand side shows the gain amplitude after all detected outliers were blanked. Blanking has maintained the smooth gain structure into the time interval that contains the strong RFI, and the entire time interval can be accurately calibrated and imaged.

4. Conclusion

The MWA employs a real-time calibration and imaging system to cope with a large data rate and fast-changing calibration parameters. While the MWA is situated in the radio-quiet Murchison Radio-astronomy Observatory, it will observe in frequency bands that are occupied by numerous communication signals. While these signals are transmitted from distant locations, they will, on

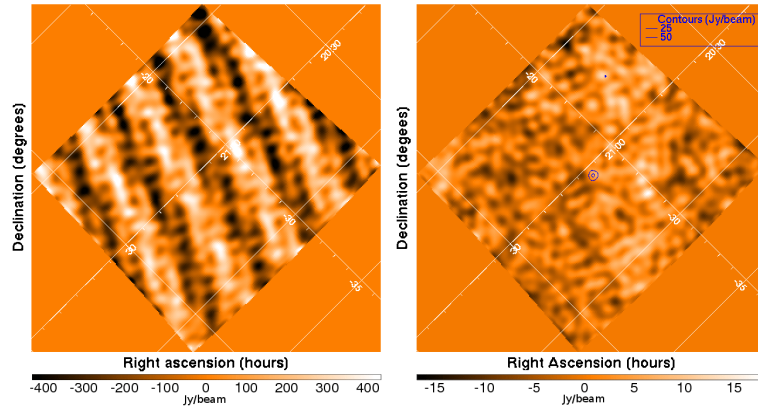


Figure 3: On the left is an image of the J2107-2526 field produced by integrating 8-second snapshots over the entire time interval without blanking. On the right is an image of the field after RFI blanking and peeling, along with contours of the unpeeled image.

occasion, reflect or refract into the receivers at levels that are orders of magnitude above the noise floor. During deep integrations the MWA real-time system will simply discard dubious data. This will require a series of data-quality tests, of which the simple median-based detector shown here will form an integral part.

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